Gallium isotopic evidence for extensive volatile loss from the Moon during its formation

Chizu Kato1,† and Frédéric Moynier1,2

The distribution and isotopic composition of volatile elements in planetary materials holds a key to the characterization of the early solar system and the Moon’s formation. The Moon and Earth are chemically and isotopically very similar. However, the Moon is highly depleted in volatile elements and the origin of this depletion is still debated. We present gallium isotopic and elemental measurements in a large set of lunar samples to constrain the origin of this volatile depletion. We show that while Ga has a geochemical behavior different from Zn, both elements show a systematic enrichment in the heavier isotopes in lunar mare basalts and Mg-suite rocks compared to the silicate Earth, pointing to a global-scale depletion event. On the other hand, the ferroan anorthosites are isotopically heterogeneous, suggesting a secondary distribution of Ga at the surface of the Moon by volatilization and condensation. The isotopic difference of Ga between Earth and the Moon and the isotopic heterogeneity of the crustal ferroan anorthosites suggest that the volatile depletion occurred following the giant impact and during the lunar magma ocean phase. These results point toward a Moon that has lost its volatile elements during a whole-scale evaporation event and that is now relatively dry compared to Earth.

INTRODUCTION

The Moon is thought to have formed from an impact between the proto-Earth and a Mars-sized impactor (1–3) or from multiple impactors (4), which resulted in a formation of a silicate vapor–rich disk (5) that later condensed to form the Moon (6). Most lunar rocks are depleted in volatile elements compared to the bulk silicate Earth (BSE) (7–9), pointing toward a volatile-depleted lunar interior, although the origin of the depletion is poorly understood (10). This general volatile depletion is more puzzling, considering that not all samples point toward a volatile-depleted Moon. For example, some olivine from the lunar pyroclastic glasses contains melt inclusions with elevated abundances of OH, F, and Cl, implying a source for these samples enriched in volatiles (11, 12). In addition, some mare basalts apatites have high OH content (13–15). However, recent measurements on plagioclase feldspar, which is a major mineral of the lunar crust, indicated that much of the lunar mantle must have <100 parts per million (ppm) of water (16).

Recent high-precision isotopic measurements of lunar samples tried to solve this problem by using Zn, Cl, Rb, and K isotopes as tracers of the origins of the volatile depletion in mare basalts (17–24). These studies have demonstrated that mare basalts are depleted in the lighter isotopes of those elements compared to the BSE, suggesting a whole-scale evaporation event such as the giant impact (20, 24), and/or during magma ocean differentiation (10, 21, 22), or during an eruptive degassing of volatiles from mare basalt lavas into a vacuum (19).

To further investigate the volatile elements in the Moon and their depletion, we analyzed, with high precision, Ga isotopic and elemental abundances on various lunar samples that were previously analyzed for their Zn isotopic composition (21). Gallium is a moderately volatile element with a 50% condensation temperature of 968 K (25), which shows a ~10-fold depletion in lunar basalts compared to their terrestrial counterparts (7). Previous studies of different crustal and magmatic terrrestrial samples demonstrated that Ga isotopes do not fractionate during magmatic processes within analytical uncertainty (26). The relatively low condensation temperature and the absence of isotopic fractionation during magmatic processes make the Ga isotope ratio a powerful tracer of high-temperature evaporation during volatile depletion events on planets.

RESULTS

Gallium concentrations and isotopic compositions (Δ266Ga = ([266Ga/69Ga]sample/[266Ga/69Ga]Ga IPGP standard − 1) × 1000) are reported in Table 1 and Fig. 1. There is a moderately good correlation between the Ga and the Zn isotopic composition of the lunar samples (Fig. 2). Because Zn (lithophile) and Ga (siderophile) do not have similar geochemical behaviors during planetary differentiation or igneous processes (Ga is incompatible, whereas Zn is fairly compatible), the correlation between Δ266Zn and Δ266Ga must be a consequence of volatile-related processes. This observation is a strong argument in favor of a volatile control on both the abundance and isotopic composition of Ga and Zn. The Ga concentrations of the different lunar samples are all below 4 ppm. Four high-Ti basalts (Δ266Ga = 0.15 to 0.57‰) and three low-Ti basalts (Δ266Ga = 0.09 to 0.32‰) demonstrate that the mare basalts are enriched in the heavier isotopes of Ga compared to the BSE [Δ266Ga = 0.00 ± 0.06‰ 2 SD; (26)] (Fig. 1). A few exceptions (two of nine samples) are isotopically light, including 14053, which coincidently also has one of the lightest Zn isotopic compositions (21). These Zn-rich samples cannot represent volatile-rich, undegassed regions of the lunar mantle because there would be no reason for these samples to be isotopically light. On the other hand, this enrichment in the lighter isotopes of Ga and Zn is best explained by contamination of the samples from the condensation of an isotopically light vapor at the surface of the samples (21) and has therefore been excluded from further discussion. Gallium concentrations in the mare basalts span a small range from 0.8 to 3.2 ppm, which is more than seven times less than their terrestrial counterparts (~20 ppm; [26]). Compared to the mare basalts, ferroan anorthosites (FANs) display larger Ga isotopic variability (Δ266Ga between −0.47 and 0.85‰; Fig. 1). In comparison, the Mg-suite rock (77215) was isotopically heavy (Δ266Ga = 0.16 ± 0.05‰ 2 SD). The pyroclastic green

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glass (15426) is isotopically lighter than the mare basalts ($d^{71}\text{Ga} = -0.35 \pm 0.05$‰). Finally, one regolith sample (78221) had an isotopic value of $d^{71}\text{Ga} = 1.16 \pm 0.05$‰ (2 SD), being the heaviest sample within this sample set. This enrichment in the heavy isotope of Ga in a regolith sample was expected on the basis of previous Zn ($\delta^{66}\text{Zn} = 1.4 \pm 0.5$‰; (21)) and K ($\delta^{41}\text{K} = 0.372 \pm 0.061$‰; (22)) isotopic studies that show that the lighter isotopes were lost during the regolith gardening by either sputtering or micrometeorite impacts.

**DISCUSSION**

The source of the mare basalts is enriched in the heavy isotope of Ga ($d^{71}\text{Ga} = 0.09$ to $0.57$‰, $n = 7$) compared to the BSE ($d^{71}\text{Ga} = 0.00 \pm 0.06$‰; (21)). Together with previous results based on Zn ($\delta^{66}\text{Zn} = 1.4 \pm 0.5$‰; (21)) and K (only two data in mare basalts: $\delta^{41}\text{K} = 0.372 \pm 0.061$‰ and $\delta^{41}\text{K} = 0.460 \pm 0.037$‰; (22)), this confirms the general enrichment of mare basalts in heavy isotopes of moderately volatile elements compared to Earth. It has previously been suggested that the heavy isotope enrichments in the mare basalts compared to the BSE in Zn were due to the fact that the source of the mare basalts was isotopically depleted by a global-scale high-temperature evaporation event (20). Isotopic enrichment and concentration depletion compared to the BSE is seen in the mare basalts for Ga, along with a variety of elements (Zn, K, S, Cl, and Rb), regardless of different sampling sites, further confirms a global-scale volatile depletion event, such as the giant impact or during evaporation from the magma ocean as the cause.

The Mg-suite sample consists of igneous cumulates that are characterized by their high Mg# together with anorthitic plagioclase, low abundance of compatible elements (for example, Ni and Co), and a small K, rare-earth elements and P (KREEP) component (29). Although the origin of the Mg-suite is debated, their plutonic origin implies that the isotopic fractionation of Ga did not occur via evaporation during degassing into a vacuum. The Ga isotopic composition of the Mg-suite sample is similar to the mare basalts and shows a slight enrichment in the heavy isotope compared to the BSE. This further confirms that the source magma was already depleted in the volatile elements, and therefore,

<table>
<thead>
<tr>
<th>Sample type</th>
<th>Sample no.</th>
<th>$d^{71}\text{Ga}$ (‰)</th>
<th>2 SD</th>
<th>$n$</th>
<th>Ga (ppm)</th>
<th>Ga (ppm)*</th>
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</thead>
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<tr>
<td>Mg-suite</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
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<td>0.16</td>
<td>0.05</td>
<td>1</td>
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<td>—</td>
</tr>
<tr>
<td>Regolith</td>
<td></td>
<td></td>
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<td></td>
<td></td>
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<tr>
<td>Soil</td>
<td>78221</td>
<td>1.16</td>
<td>0.05</td>
<td>1</td>
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<td>—</td>
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<td>Pyroclastic glass</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
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<tr>
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<td>−0.35</td>
<td>0.05</td>
<td>1</td>
<td>1.3</td>
<td>4.7</td>
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<td>FANs</td>
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<td></td>
<td></td>
<td></td>
<td></td>
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<tr>
<td>Anorthosite (with melt)</td>
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<td>1</td>
<td>0.5</td>
<td>—</td>
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<tr>
<td>Cataclastic anorthosite</td>
<td>60015</td>
<td>0.17</td>
<td>0.05</td>
<td>1</td>
<td>1.4</td>
<td>—</td>
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<td>Ferroan anorthosite</td>
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<td>1</td>
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<td>3.1</td>
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<td>0.05</td>
<td>1</td>
<td>1.9</td>
<td>3.5–4.2</td>
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<td></td>
<td></td>
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<td>Basalt (low Ti)</td>
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<td>0.05</td>
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<td>3.2</td>
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<td>Olivine basalt</td>
<td>12012</td>
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<td>3</td>
<td>1.1</td>
<td>4.8</td>
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<td>Olivine basalt</td>
<td>12040</td>
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<td>0.05</td>
<td>3</td>
<td>1.4</td>
<td>1.9</td>
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<td>Mare basalt (Al-rich)</td>
<td>14053</td>
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<td>0.05</td>
<td>1</td>
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<td>2.5–4.8</td>
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<td></td>
<td></td>
</tr>
<tr>
<td>Basalt (high Ti)</td>
<td>70017</td>
<td>0.15</td>
<td>0.01</td>
<td>3</td>
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<td>3.1–21</td>
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<td>0.06</td>
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<td>4.0–4.7</td>
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<td>Basalt (high Ti)</td>
<td>70135</td>
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<td>0.02</td>
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<td>7–16</td>
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<tr>
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<td>0.03</td>
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<td>—</td>
<td>4.8</td>
</tr>
<tr>
<td>Ilmenite basalt (high Mg/Fe)</td>
<td>12005</td>
<td>0.22</td>
<td>0.01</td>
<td>3</td>
<td>1.1</td>
<td>—</td>
</tr>
<tr>
<td>BSE</td>
<td></td>
<td>0.00†</td>
<td>0.05</td>
<td>3</td>
<td>3.9§</td>
<td>—</td>
</tr>
</tbody>
</table>

*Data from Meyer (47). †Used the 2 SD estimated from Kato et al. (26). ‡Data from Kato et al. (26). §Data from McDonough (48).
the heavy Ga, Zn, K, and Cl isotopic compositions of mare basalts are not due to degassing of their lava in a vacuum.

The lunar pyroclastic green glass (15426) is highly enriched in volatile elements (including Ga and Zn) at the surface compared to the interior (30–32), which is also enriched in the lighter isotopes of Ga and Zn ($\delta^{71} \text{Ga} = -0.35\%$, $\delta^{66} \text{Zn} = -0.98\%$). These results further confirm that they were formed from lunar fire fountaining from the lunar interior (33, 34). During the ascent of the magma, the pressure decrease induces the formation of gas bubbles enriched in the lighter isotopes of Ga and Zn due to kinetic fractionation. This volatile-rich and isotopically light gas then condenses onto the surface of the quenched glass after eruption.

FANs are considered to have been produced from the crystallization of the lunar magma ocean as flotation lunar crust (35). They would therefore represent the primary crust of the Moon and are the oldest rock samples of the Moon (36). The isotopic variation in a single sample set is the largest seen within the samples analyzed in this study, with a maximum range of 1.3% between the brecciated 62255 ($\delta^{71} \text{Ga} = -0.47\%$) and unbrecciated pristine 15415 ($\delta^{71} \text{Ga} = 0.85\%$). Because Ga isotopes do not fractionate (within analytical uncertainty) during fractional crystallization (26), two scenarios can account for the Ga isotopic variability recorded in FANs: (i) the isotopic variations are inherited from the heterogeneity of the original source magma, or (ii) secondary redistribution of Ga at the surface of the Moon fractionated Ga isotopes. Gallium is moderately incompatible and behaves similarly to Al during magmatic differentiation due to their comparable ionic radii, valence, and ionization potential, and therefore, the Ga/Al ratios do not fractionate during the last stages of crystallization or by partial melting (37). The FANs exhibit variable Ga/Al ratios (Table 2), which suggests that the difference in the Ga/Al ratio is rather due to local volatilization and condensation. Impact on the surface of the FANs will selectively evaporate the light isotope, and later, the isotopically light gas will condense onto another surface, forming the isotopically heterogeneous FANs. In addition, some FANs may represent mixed lithologies (38), and therefore, it is possible that the observed trend represents secondary mixing between lithologies with variable isotopic compositions. Zinc and Ga isotope ratios follow systematic variations between FAN samples (Fig. 2), further suggesting that high-temperature evaporation events lead to the origin of these variations (21). In addition, the Fe/Zn ratio, which also does not fractionate during partial melting under reducing conditions (39), can be used as a tracer of volatile depletion because Zn is more volatile than Fe (25). Following this logic, the low Fe/Zn ratio of lunar mare basalts compared to terrestrial basalts was used as an evidence for the volatile depletion of the Moon (40). Similarly, the variations between FAN samples analyzed here suggest volatilization as the cause for the isotopic fractionation. This interpretation is confirmed by the texture of the samples; the unbrecciated pristine 15415 (41) is enriched in heavy isotopes for both Ga ($\delta^{71} \text{Ga} = 0.85\%$) and Zn ($\delta^{66} \text{Zn} = 4.24\%$), whereas the brecciated 62255 ($\delta^{71} \text{Ga} = -0.47\%$; $\delta^{66} \text{Zn} = 0.00\%$) incorporated the isotopically light Ga- and Zn-rich vapor during cataclasis or exposure on the lunar surface. This concludes that the FANs do not represent the isotopic composition of their source magma, and this composition is controlled by billions of years of impacts on the surface of the Moon.

Mare basalts and the plutonic Mg-suite sample are enriched in the heavy isotopes of Ga and Zn compared to the BSE. This supports a magma source originally depleted in volatile elements and not volatile loss during degassing into a vacuum. Because all chondrites are isotopically lighter in Ga isotopes than the BSE (and therefore lighter than the Moon) (42), this heavy isotope enrichment cannot be due to the composition of the impactor. In turn, it would require a whole-scale evaporation event for the loss of Ga, either by the giant impact or during

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**Table 2. Ga/Al ratio of the FANs.**

<table>
<thead>
<tr>
<th>Sample type</th>
<th>Sample no.</th>
<th>Ga (ppm)</th>
<th>Al (ppm)*</th>
<th>Ga/Al*10^4 (ppm/ppm)</th>
</tr>
</thead>
<tbody>
<tr>
<td>FANs</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>Anorthosite (with melt)</td>
<td>62255</td>
<td>0.5</td>
<td>160786</td>
<td>0.03</td>
</tr>
<tr>
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<td>60015</td>
<td>1.4</td>
<td>185776</td>
<td>0.08</td>
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<tr>
<td>Ferroan anorthosite</td>
<td>15415</td>
<td>1.1</td>
<td>188440</td>
<td>0.06</td>
</tr>
<tr>
<td>Noritic anorthosite</td>
<td>67955</td>
<td>1.9</td>
<td>143674</td>
<td>0.13</td>
</tr>
</tbody>
</table>

*Data from Meyer (47).
the lunar magma ocean. Outgassing during planetary impacts and/or magma ocean phases plays a crucial role in setting the volatile inventories of planets. For the Earth-Moon system, it has been shown that hydrodynamic escape from the lunar disk is not possible and that volatiles (including water) cannot be lost to space (43). On the other hand, dynamical biphased vapor-melt numerical modeling of the proto-lunar disk shows that because of magnetorotational instability, the vapor phase is viscous and accretes preferentially to Earth, leaving volatile-depleted material to form the Moon (6). Furthermore, the volatile loss during a magma ocean phase of the Moon would predict that this effect would be even more pronounced on low mass bodies. This is observed in the angrite and HED (howardite-eucrite-diogenite) parent bodies, which are both highly depleted in volatile elements. Therefore, if the volatile loss occurred during the magma ocean phase, it would be expected that Ga isotopes would be highly fractionated in HED and angrite meteorites. Regardless of which whole-scale evaporation event occurred, our results imply that planets accreted principally from highly differentiated planetesimals and must have accreted volatile depleted.

**REFERENCES AND NOTES**


36. E. L. Connelly, J. N. Connelly, M. Boyet, R. W. Carlson, Chronological evidence that the Moon is either young or did not have a global magma ocean. Nature 477, 70–72 (2011).


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